

# Astronomical Notes

## Astronomische Nachrichten

Founded by H. C. Schumacher in 1821

### Editors

K. G. Strassmeier (Potsdam/Editor-in-Chief),  
G. Hasinger (Garching), R.-P. Kudritzki (Honolulu),  
T. Montmerle (Grenoble), H. W. Yorke (Pasadena)

 **WILEY-VCH**

**REPRINT**

# Large-scale magnetic flux concentrations from turbulent stresses

A. Brandenburg<sup>1,2,\*</sup>, N. Kleeorin<sup>3</sup>, and I. Rogachevskii<sup>3</sup>

<sup>1</sup> NORDITA, AlbaNova University Center, Roslagstullsbacken 23, SE 10691 Stockholm, Sweden

<sup>2</sup> Department of Astronomy, AlbaNova University Center, Stockholm University, SE 10691 Stockholm, Sweden

<sup>3</sup> Department of Mechanical Engineering, The Ben-Gurion University of the Negev, POB 653, Beer-Sheva 84105, Israel

Received 2009 Oct 12, accepted 2009 Nov 20

Published online 2009 Dec 30

**Key words** instabilities – magnetic fields – magnetohydrodynamics (MHD) – turbulence

In this study we provide the first numerical demonstration of the effects of turbulence on the mean Lorentz force and the resulting formation of large-scale magnetic structures. Using three-dimensional direct numerical simulations (DNS) of forced turbulence we show that an imposed mean magnetic field leads to a decrease of the turbulent hydromagnetic pressure and tension. This phenomenon is quantified by determining the relevant functions that relate the sum of the turbulent Reynolds and Maxwell stresses with the Maxwell stress of the mean magnetic field. Using such a parameterization, we show by means of two-dimensional and three-dimensional mean-field numerical modelling that an isentropic density stratified layer becomes unstable in the presence of a uniform imposed magnetic field. This large-scale instability results in the formation of loop-like magnetic structures which are concentrated at the top of the stratified layer. In three dimensions these structures resemble the appearance of bipolar magnetic regions in the Sun. The results of DNS and mean-field numerical modelling are in good agreement with theoretical predictions. We discuss our model in the context of a distributed solar dynamo where active regions and sunspots might be rather shallow phenomena.

© 2010 WILEY-VCH Verlag GmbH & Co. KGaA, Weinheim

## 1 Introduction

Turbulence effects generally refer to the occurrence of correlations between components of velocity, temperature, and/or magnetic fields at small scales. A typical example is turbulent viscosity, which results from the spatial exchange of turbulent eddies characterized by velocity correlations. This leads to the dissipation of energy at small scales.

However, there is also the possibility of additional (e.g. non-diffusive) turbulence effects, as is perhaps best known in mean-field electrodynamics and dynamo theory. Here one models the effects of the mean electromotive force, i.e. the turbulence effects of velocity and magnetic field fluctuations, on the evolution of the mean field. This can lead to the occurrence of the famous  $\alpha$  effect, in addition to turbulent magnetic diffusion, turbulent diamagnetic velocity, etc. (Moffatt 1978; Krause & Rädler 1980). Another example is the  $\Lambda$  effect in rotating anisotropic hydrodynamic turbulence, which can lead to the occurrence of differential rotation in cosmic bodies such as the Sun (Rüdiger 1980, 1989; Rüdiger & Hollerbach 2004). In that case the relevant correlations come from the mean Reynolds stress tensor and its dependence on the local angular velocity.

A related example is the combined Reynolds and Maxwell turbulent stress tensor and its dependence on the mean magnetic field. The first analytic calculations of the dependence of the turbulent Reynolds stress on the mean magnetic field in the framework of the first-order smooth-

ing approximation were performed by Rädler (1974) and Rüdiger (1974). Later, also the combined effects of the turbulent Reynolds and Maxwell stress tensors were considered (Kleeorin et al. 1989, 1990; Rüdiger & Kichatinov 1990). It was noticed that this can lead to a local reduction of the total turbulent pressure and hence to the possibility of self-induced concentrations of large-scale magnetic fields (Kleeorin et al. 1989, 1990, 1996; Kleeorin & Rogachevskii 1994; Rogachevskii & Kleeorin 2007). Such a process may play an important role in the formation of sunspots and active regions in the Sun. It may be complementary to the magnetically induced suppression of the turbulent energy flux, which would lead to further cooling and hence a further concentration of the structures (Kitchatinov & Mazur 2000).

The currently leading explanation for the formation of sunspots is related to the emergence of deeply rooted magnetic flux tubes (Parker 1955, 1982, 1984). Such flux tubes are generally believed to be produced and ‘stored’ near the bottom of the convection zone (Spiegel & Weiss 1980). The storage of magnetic fields and the formation of flux tubes in the overshoot layer near the bottom of the solar convective zone was investigated in a number of publications (see, e.g., Spruit 1981; Spruit & van Ballegoijen 1982; Schüssler et al. 1994; Moreno-Inertis et al. 1996; Tobias et al. 2001; Tobias & Hughes 2004). However, in order that the tubes retain their basic east–west orientation during their ascent over many pressure scale heights, the magnetic field must be strong enough (Choudhuri & D’Silva 1990) and is estimated to be around  $10^5$  G at the bottom of the convection zone

\* Corresponding author: brandenb@nordita.org

(D'Silva & Choudhuri 1993). Such fields would be up to a hundred times stronger than the equipartition value, which is one of several arguments that have led to the idea that flux emergence of dynamo-generated fields might instead be a shallow phenomenon (Brandenburg 2005; Schatten 2009). Such a scenario appears also compatible with solar subsurface flows, as inferred from local helioseismology (Zhao et al. 2001; Kosovichev 2002). In particular, Zhao et al. (2004) and Hindman et al. (2009) find the presence of converging flows around active regions at radii as large as 100–200 Mm. It appears that these convergent flows might actually be the source of the formation of active regions and perhaps sunspots rather than a consequence, as is normally believed (Parker 1979a; Hurlburt & Rucklidge 2000). Of course, in the immediate proximity of individual spots one observes outgoing flows. Those are probably superficial, less than 1–2 Mm deep, and appear to be due to the dynamical effects of magnetoconvection in an inclined magnetic field of the penumbra (e.g., Thomas et al. 2002; Heinemann et al. 2007; Rempel et al. 2009; Kitiashvili et al. 2009).

The goal of this paper is to investigate the effects of turbulence on the mean Lorentz force by means of direct numerical simulations (DNS) for forced turbulence and to study the instability of a uniform large-scale magnetic field in an adiabatically stratified layer by means of mean-field numerical modelling based on parameterizations both of analytic formulae by Rogachevskii & Kleeorin (2007) and the results of DNS for forced turbulence. In order to study the essence of the effect, we make several simplifications by neglecting the energy equation, i.e. the specific entropy is assumed to be strictly constant in space and time. In the mean-field numerical modelling we neglect the suppression of turbulent magnetic diffusivity and turbulent viscosity, and omit correlations with density fluctuations. Nevertheless, the mean density is allowed to evolve fully self-consistently according to the usual continuity equation.

## 2 Turbulence effects on mean Lorentz force

Throughout this paper we adopt units for the magnetic field where the vacuum permeability is equal to unity, i.e. the magnetic pressure is given by  $\frac{1}{2}\mathbf{B}^2$ .

### 2.1 General considerations

We use the equations of mean-field magnetohydrodynamics (MHD). These equations are obtained by averaging the original MHD equations over small-scale fluctuations. This technique is best known in the case of the induction equation (Moffatt 1978; Krause & Rädler 1980). In this study we are mainly interested in effects of turbulence on the mean Lorentz force. Let us consider the momentum equation,

$$\frac{\partial}{\partial t}\rho\mathbf{U} = -\frac{\partial}{\partial x_j}\Pi_{ij}, \quad (1)$$

where

$$\Pi_{ij} = \rho U_i U_j + \delta_{ij} \left( p + \frac{1}{2}\mathbf{B}^2 \right) - B_i B_j - \sigma_{ij} \quad (2)$$

is the momentum stress tensor,  $\mathbf{U}$  and  $\mathbf{B}$  are velocity and magnetic fields,  $p$  and  $\rho$  are the fluid pressure and density,  $\delta_{ij}$  is the unit Kronecker tensor,  $\sigma_{ij} = 2\rho\nu S_{ij}$  is the viscous stress tensor, with

$$S_{ij} = \frac{1}{2}(U_{i,j} + U_{j,i}) - \frac{1}{3}\delta_{ij}\nabla \cdot \mathbf{U} \quad (3)$$

being the traceless rate of the strain tensor, and  $\nu$  is the kinematic viscosity. Ignoring the turbulent correlations with density fluctuations for low-Mach number turbulence, the averaged momentum equation is

$$\frac{\partial}{\partial t}\bar{\rho}\bar{\mathbf{U}} = -\frac{\partial}{\partial x_j}\bar{\Pi}_{ij}, \quad (4)$$

where  $\bar{\Pi}_{ij} = \bar{\Pi}_{ij}^m + \bar{\Pi}_{ij}^f$  is the mean momentum stress tensor split into contributions resulting entirely from the mean field (indicated by superscript m) and those of the fluctuating field (indicated by superscript f). The tensor  $\bar{\Pi}_{ij}^m$  has the same form as Eq. (2), but all quantities have now attained an overbar, i.e.

$$\bar{\Pi}_{ij}^m = \bar{\rho}\bar{U}_i\bar{U}_j + \delta_{ij} \left( \bar{p} + \frac{1}{2}\bar{\mathbf{B}}^2 \right) - \bar{B}_i\bar{B}_j - \bar{\sigma}_{ij}. \quad (5)$$

We emphasize here that  $\bar{p}$  is just the mean gas pressure and  $\bar{\sigma}_{ij}$  is the average of the microscopic viscous stress tensor,  $\sigma_{ij}$ . In this paper we consider isotropic background turbulence (i.e., turbulence with a zero-mean magnetic field). The contribution from the fluctuating fields, in turn, is split into contributions that are independent of the mean fields (and hence isotropic and proportional to  $\delta_{ij}$ ) and contributions which depend on the mean fields,

$$\bar{\Pi}_{ij}^f = p_{t0}\delta_{ij} - \sigma_{ij}^{\text{eff}}. \quad (6)$$

Here,  $p_{t0}$  is the turbulent pressure in the absence of a mean magnetic field and  $\sigma_{ij}^{\text{eff}} = \sigma_{ij}^K + \sigma_{ij}^M$  quantifies the turbulent viscosity,  $\sigma_{ij}^K = 2\bar{\rho}\nu_t\bar{S}_{ij}$ , and the additional effects of the mean magnetic field on the effective stress tensor. The combination  $p_{t0}\delta_{ij} - \sigma_{ij}^M$  results from the fluctuations of velocity and magnetic fields,  $\mathbf{u} = \mathbf{U} - \bar{\mathbf{U}}$  and  $\mathbf{b} = \mathbf{B} - \bar{\mathbf{B}}$ , respectively, and is given by

$$p_{t0}\delta_{ij} - \sigma_{ij}^M = \bar{\rho}\overline{u_i u_j} + \frac{1}{2}\delta_{ij}\overline{\mathbf{b}^2} - \overline{b_i b_j}. \quad (7)$$

In the absence of a mean magnetic field, the turbulent background pressure is

$$p_{t0} = \frac{1}{6}\overline{\mathbf{b}^2} + \frac{1}{3}\rho\overline{\mathbf{u}^2}, \quad (8)$$

where the subscripts 0 on  $\mathbf{b}^2$  and  $\mathbf{u}^2$  indicate values in the absence of the mean magnetic field. Magnetic fluctuations  $\mathbf{b}$  are generated both by small-scale dynamo action and by tangling of the mean magnetic field by velocity fluctuations. On the other hand, the velocity fluctuations also have two contributions, those that depend on the mean magnetic field and those that do not.

Following Rogachevskii & Kleeorin (2007), the part of the effective stress tensor that depends on the mean magnetic field is parameterized as

$$\sigma_{ij}^M(\bar{\mathbf{B}}) = -q_s\bar{B}_i\bar{B}_j + \frac{1}{2}\delta_{ij}q_p\bar{\mathbf{B}}^2, \quad (9)$$

where  $q_s$  and  $q_p$  are functions of the mean field,  $\bar{\mathbf{B}}$ , and the magnetic Reynolds number,  $R_m$ . Equation (9) implies that

the effective mean Lorentz force that takes into account the effects of turbulence, can be written as

$$\overline{\rho \mathcal{F}}_M = -\frac{1}{2} \nabla [(1 - q_p) \overline{\mathbf{B}}^2] + \overline{\mathbf{B}} \cdot \nabla [(1 - q_s) \overline{\mathbf{B}}]. \quad (10)$$

The detailed analytic expressions for  $q_s(\overline{\mathbf{B}})$  and  $q_p(\overline{\mathbf{B}})$  have been given by Rogachevskii & Kleeorin (2007). Asymptotic formulae for the nonlinear functions,  $q_p(\overline{\mathbf{B}})$  and  $q_s(\overline{\mathbf{B}})$ , are given below. For this purpose we define  $\beta \equiv \overline{B}/B_{\text{eq}}$ , where  $B_{\text{eq}} = (\rho \overline{u_0^2})^{1/2}$  is the equipartition field strength.

For very weak mean magnetic fields,  $4\beta \ll \text{Rm}^{-1/4}$ ,  $q_p$  and  $q_s$  are approximately constant and given by

$$q_p(\beta) = \frac{4}{45} (1 + 9 \ln \text{Rm}),$$

$$q_s(\beta) = \frac{2}{15} (1 + 4 \ln \text{Rm});$$

for  $\text{Rm}^{-1/4} \ll 4\beta \ll 1$  we have

$$q_p(\beta) = \frac{16}{25} [1 + 5 |\ln(4\beta)| + 32 \beta^2],$$

$$q_s(\beta) = \frac{32}{15} \left[ |\ln(4\beta)| + \frac{1}{30} + 12 \beta^2 \right],$$

and for strong fields,  $4\beta \gg 1$ , we have, according to Rogachevskii & Kleeorin (2007),

$$q_p(\beta) = 1/6\beta^2, \quad q_s(\beta) = \pi/48\beta^3.$$

In Sect. 2.2 we present DNS evidence that the functions  $q_p(\overline{\mathbf{B}})$  and  $q_s(\overline{\mathbf{B}})$  are positive, indicating the possibility of a reduction of the effective Lorentz force, i.e., a decrease of the effective magnetic pressure and magnetic tension in small-scale turbulence.

## 2.2 DNS of turbulence effects on mean Lorentz force

In order to study turbulence effects on the mean Lorentz force and to determine the functions  $q_p(\overline{\mathbf{B}})$  and  $q_s(\overline{\mathbf{B}})$  from DNS, we consider forced turbulence in a periodic three-dimensional domain in the presence of an imposed uniform magnetic field, say  $\overline{\mathbf{B}}_0 = (\overline{B}_0, 0, 0)$ . We determine  $q_s$  and  $q_p$  from Eqs. (7) and (9) for  $i = j = x$ ,

$$p_{t0} + \frac{1}{2} (\overline{b_x^2} - \overline{b_y^2} - \overline{b_z^2}) - \overline{\rho u_x^2} = (\frac{1}{2} q_p - q_s) \overline{B}_0^2, \quad (11)$$

and  $i = j = y$ ,

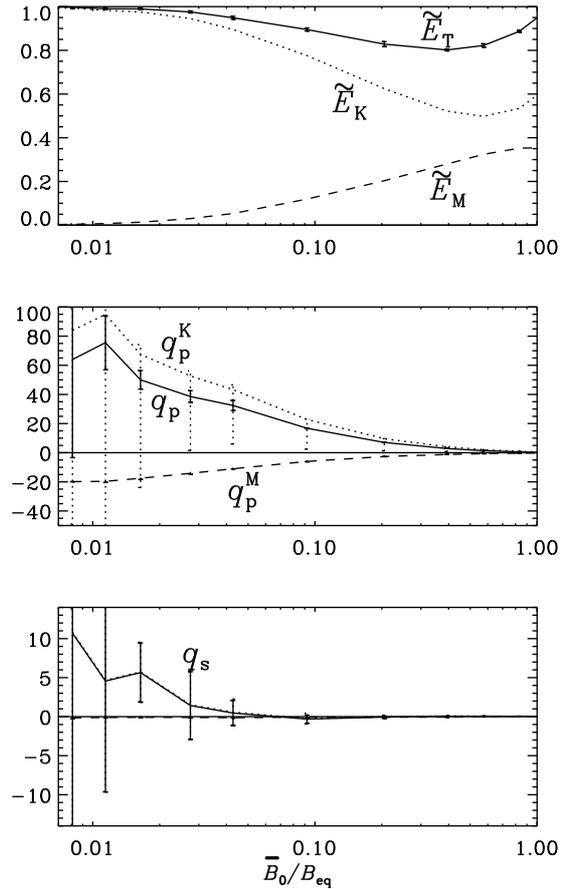
$$p_{t0} + \frac{1}{2} (\overline{b_y^2} - \overline{b_x^2} - \overline{b_z^2}) - \overline{\rho u_y^2} = \frac{1}{2} q_p \overline{B}_0^2, \quad (12)$$

where  $p_{t0}$  is given by Eq. (8). First, we determine  $p_{t0}$  from a simulation with  $\overline{B}_0 = 0$ . Then, we use Eq. (12) to determine  $q_p(\overline{\mathbf{B}})$ . Finally, to determine  $q_s(\overline{\mathbf{B}})$  we subtract Eq. (11) from Eq. (12), i.e.

$$(\overline{b_y^2} - \overline{b_x^2}) - \overline{\rho (u_y^2 - u_x^2)} = q_s \overline{B}_0^2. \quad (13)$$

In order to determine separately the effects of the mean field on the turbulent Maxwell and Reynolds stresses we also compute their respective contributions  $q_p^M + q_p^K = q_p$  and  $q_s^M + q_s^K = q_s$ . We consider here cases where  $1 \ll \text{Rm} < \text{Rm}_c$ , i.e.,  $\text{Rm}$  is smaller than the threshold  $\text{Rm}_c$  for small-scale dynamo action. Thus, we have

$$q_p^M = (\overline{b_y^2} - \overline{b_x^2} - \overline{b_z^2}) / \overline{B}_0^2, \quad q_p^K = 2(p_{t0} - \overline{\rho u_y^2}) / \overline{B}_0^2. \quad (14)$$



**Fig. 1**  $\overline{B}_0$  dependence of the normalized turbulent energy  $\tilde{E}_T = E_T/E_0$ , where  $E_0$  is the value of  $E_T$  for  $\overline{B}_0 = 0$ , together with the contributions from kinetic and magnetic energies,  $\tilde{E}_K = E_K/E_0$  and  $\tilde{E}_M = E_M/E_0$  (upper panel), as well as the coefficients  $q_p$ ,  $q_p^K$ ,  $q_p^M$ , and  $q_s$  (second and third panels), obtained from DNS for  $\text{Re} = 180$  and  $\text{Rm} = 45$ .

In Fig. 1 we show the results from the simulations of forced turbulence in the presence of an imposed magnetic field. Here we have also plotted the total turbulent energy,  $E_T = E_K + E_M$ , where

$$E_K = \overline{\rho u^2}/2, \quad E_M = \overline{\mathbf{b}^2}/2 \quad (15)$$

are the energy densities of velocity and magnetic fluctuations, respectively. In all cases we adopt volume averages over a triply-periodic domain. The simulations were performed with the PENCIL CODE<sup>1</sup>, which uses sixth-order explicit finite differences in space and third-order accurate time stepping method. They are similar to the DNS of Haugen & Brandenburg (2004), which also have an imposed uniform field, but the value of the magnetic Reynolds number was large enough to support dynamo action. As in earlier work, the forcing function consists of plane non-polarized waves with an average wavenumber  $k_f = 1.5 k_1$ , where  $k_1 = 2\pi/L$  is the smallest wavenumber that fits into a 3D periodic domain of size  $L$ . We use an isothermal

<sup>1</sup> <http://pencil-code.googlecode.com>

equation of state with constant sound speed  $c_s$ . The forcing strength is arranged such that the turbulent rms velocity,  $u_{\text{rms}}$ , is below  $c_s$ . In all our runs the maximum turbulent Mach number,  $u_{\text{rms}}/c_s$ , is around 0.2, so that compressibility effects are still weak.

The fluid Reynolds number based on the forcing wavenumber,  $k_f = 2\pi/l_f$ , is  $\text{Re} = u_{\text{rms}}/\nu k_f \approx 180$  and the magnetic Prandtl number is  $\nu/\eta = 0.25$ , so the magnetic Reynolds number,  $R_m = u_{\text{rms}}/\eta k_f \approx 45$ , is just small enough so that no small-scale dynamo is excited.<sup>2</sup> Note that the magnetic Reynolds number,  $\text{Rm}$ , used in Sect. 2.1 is based on forcing scale  $l_f = 2\pi/k_f$ , and thus  $\text{Rm} = 2\pi R_m$ . In all cases we normalize the field strength in terms of the equipartition value,  $B_{\text{eq}} = \rho_0^{1/2} u_{\text{rms}}$ . The number of mesh points is  $64^3$ .

As follows from our DNS study, both functions  $q_p(\overline{B})$  and  $q_s(\overline{B})$  are positive and exceed unity for weak fields. The error bars are obtained by calculating averages over each of three equally long intervals of the full time series and taking the largest deviation from the full averages. For small values of  $B_0/B_{\text{eq}}$ , the turbulent fluctuations dominate over the effects of shredding the imposed field, which increases the error bars. Indeed, for zero mean field the relative anisotropy,  $(2p_{t0} - \overline{p}u_x^2)/3p_{t0}$ , is about  $\pm 10^{-3}$ . In order for the absolute error of  $q_p$  to stay below 10, we must have  $B_0/B_{\text{eq}} > 0.01$ . For those field strengths the results are in agreement with the prediction of Rogachevskii & Kleeorin (2007). Note that both theory and simulations suggest that  $q_p > 2q_s$ .

### 2.3 Physics of turbulence effects on Lorentz force

The physics of the effect of turbulence on the large-scale Lorentz force is as follows. The equation of state for isotropic turbulence is given by

$$p_{t0} = \frac{1}{3}E_M + \frac{2}{3}E_K \quad (16)$$

(see Eq. 8 and also, e.g., Landau & Lifshitz 1975, 1984), where  $E_K$  and  $E_M$  were defined by Eq. (15). The total energy density  $E_T = E_K + E_M$  of homogeneous turbulence with a mean magnetic field  $\overline{B}$  is determined by the equation

$$\frac{\partial E_T}{\partial t} = I_T - \frac{E_T}{\tau_0} + \eta_t (\nabla \times \overline{B})^2, \quad (17)$$

where  $\tau_0$  is the correlation time of the turbulent velocity field in the maximum scale  $l_f$  of turbulent motions,  $I_T$  is the energy source of turbulence,  $\eta_t$  is the turbulent magnetic diffusion. For a given time-independent source of turbulence  $I_T$  and for  $t \gg \tau_0$  the total energy density of the turbulence reaches a steady state

$$E_T = \text{const} = \tau_0 I_T, \quad (18)$$

where we neglect a small magnetic source  $\eta_t (\nabla \times \overline{B})^2$  of the turbulence [that is of the order of  $O(l_f^2/H_B^2)$ ], and  $H_B$  is the characteristic scale of the mean magnetic field spatial

<sup>2</sup> Following Haugen et al. (2004) and Schekochihin et al. (2005, 2007), the critical value of  $R_m$  is between 70 and 80 for this value of the magnetic Prandtl number.

variations]. The approximate constancy of  $E_T$  is compatible with DNS, where we found only a small decrease (20%) for strong (equipartition strength) mean fields (see the upper panel of Fig. 1). The reason for the departure is possibly a dependence of  $\tau_0$  on  $\overline{B}$ . However, this decrease only enhances the modification of the Lorentz force by turbulence.

Equations (16) and (18) imply that the change of turbulent pressure  $\delta p_{t0}$  is proportional to the change of the magnetic energy density  $\delta E_M$ . In particular, using  $\delta E_K = -\delta E_M$ , one expects  $\delta p_{t0} = -(1/3)\delta E_M$ . Therefore, the total turbulent pressure is reduced when magnetic fluctuations are generated.

For a non-zero large-scale mean magnetic field  $\overline{B}$ , the change of the magnetic energy density  $\delta E_M$  is proportional to  $\overline{B}^2$ . Therefore, the total turbulent pressure is given by  $P_t = p_{t0} - \frac{1}{2}q_p \overline{B}^2$ , where  $p_{t0}$  is the turbulent pressure in a flow with a zero mean magnetic field. When magnetic fluctuations are generated, we have  $q_p > 0$ . Now we examine the part,  $P_{\text{eff}}(\overline{B})$ , of the total pressure,  $P_{\text{tot}} = \overline{p} + P_t + \frac{1}{2}\overline{B}^2$ , that depends on the mean magnetic field,  $\overline{B}$ , i.e., we consider the effective mean magnetic pressure that takes into account the contribution of turbulence,  $P_{\text{eff}}(\overline{B}) = \frac{1}{2}(1 - q_p)\overline{B}^2$ . We study the case when  $\overline{p} \gg \overline{B}^2/2$ , and therefore, the total pressure  $P_{\text{tot}}$  is always positive, while the effective mean magnetic pressure  $P_{\text{eff}}(\overline{B})$  may be negative when  $q_p > 1$ .

The modification of the mean Lorentz force can result in the excitation of a long wavelength instability even in an initially uniform mean magnetic field in a density stratified layer (Kleeorin et al. 1990, 1996; Rogachevskii & Kleeorin 2007). Indeed, the growth rate of the instability for the perturbations perpendicular to both the gravity field  $\mathbf{g}$  and the mean magnetic field  $\overline{B}$  is given by

$$\lambda = \frac{c_a}{H_\rho} \sqrt{(1 - q_p) \left( \frac{H_\rho}{H_B} - 1 \right)}, \quad (19)$$

where  $c_a = \overline{B}/\rho_0^{1/2}$  is the Alfvén speed,  $H_\rho^{-1} = |\nabla \ln \rho_0|$ ,  $H_B^{-1} = |\nabla \ln \overline{B}|$ , and we have, for simplicity, neglected the dissipation processes due to turbulent viscosity and turbulent magnetic diffusion. For  $q_p > 1$  the instability is excited when  $H_\rho < H_B$ , i.e., it occurs even in an initially uniform mean magnetic field.

The mechanism of the instability can be understood as follows. An isolated magnetic tube moving upward becomes lighter than the surrounding plasma since the decrease of the magnetic field in it, due to expansion of the tube, is accompanied with an increase of the magnetic pressure inside the tube. This increase, due to the fact that the effective magnetic pressure is negative, leads to a decrease of the density inside the tube and to the appearance of a buoyancy force. It results in the further upward displacement of the flux tube, i.e. it causes the excitation of the instability. The instability causes the formation of inhomogeneities and thus magnetic structures. The energy for this instability is supplied by the small-scale turbulence. In contrast, the free

energy in Parker's magnetic buoyancy instability, that is excited when  $H_\rho > H_B$ , is drawn from the gravitational field (Newcomb 1961; Parker 1966). The growth rate of Parker's magnetic buoyancy instability is determined by Eq. (19) for  $q_p = 0$ .

Magnetic buoyancy in astrophysics applies usually to two different situations (see, e.g., Priest 1982). The first corresponds to a problem described by Parker (1966, 1979b) and Gilman (1970a, 1970b). They considered the instability of a stratified continuous magnetic field and did not invoke a magnetic flux tube. The other situation was considered by Parker (1955), Spruit (1981), Spruit & van Ballegoijen (1982), Ferriz-Mas & Schüssler (1993), and Schüssler et al. (1994), who studied buoyancy of horizontal magnetic flux tubes.

In the present study we investigate the first situation, i.e., we study the large-scale instability of a continuous (diffusive) magnetic field in small-scale turbulence. This instability is caused by turbulent velocity and magnetic fluctuations and leads to the formation of large-scale magnetic structures (see Sect. 3).

### 3 Mean-field numerical modelling

In order to understand in more detail the appearance of magnetic structures from this large-scale instability we consider now numerical solutions of the mean-field MHD equations in a density stratified layer. We apply a mean-field model that includes for the first time the effects of turbulence on the mean Lorentz force. We solve the evolution equations for mean velocity  $\overline{\mathbf{U}}$ , mean density  $\overline{\rho}$ , and mean vector potential  $\overline{\mathbf{A}}$  in the form

$$\frac{\partial \overline{\mathbf{U}}}{\partial t} = -\overline{\mathbf{U}} \cdot \nabla \overline{\mathbf{U}} - c_s^2 \nabla \ln \overline{\rho} + \mathbf{g} + \overline{\mathcal{F}}_M + \overline{\mathcal{F}}_{K,\text{tot}}, \quad (20)$$

$$\frac{\partial \ln \overline{\rho}}{\partial t} = -\overline{\mathbf{U}} \cdot \nabla \ln \overline{\rho} - \nabla \cdot \overline{\mathbf{U}}, \quad (21)$$

$$\frac{\partial \overline{\mathbf{A}}}{\partial t} = \overline{\mathbf{U}} \times \overline{\mathbf{B}} - (\eta_t + \eta) \overline{\mathbf{J}}, \quad (22)$$

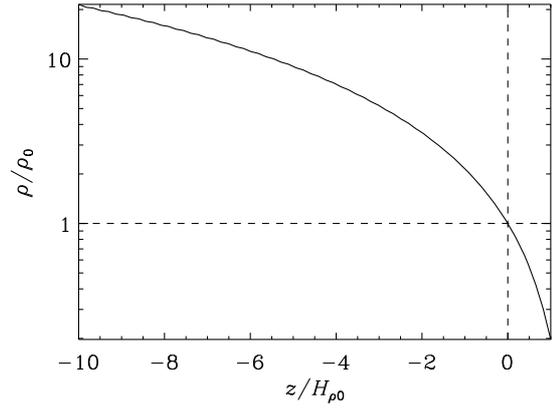
where  $\overline{\mathcal{F}}_M$  is given by Eq. (10), and  $\overline{\mathcal{F}}_{K,\text{tot}} = \overline{\mathcal{F}}_K + \overline{\mathcal{F}}_{\text{visc}}$  with  $\overline{\mathcal{F}}_K = \nabla \cdot \overline{\boldsymbol{\sigma}}^K$  and with  $\overline{\mathcal{F}}_{\text{visc}} = \nabla \cdot \overline{\boldsymbol{\sigma}}$ , so that

$$\overline{\mathcal{F}}_{K,\text{tot}} = (\nu_t + \nu) (\nabla^2 \overline{\mathbf{U}} + \nabla \nabla \cdot \overline{\mathbf{U}} + 2\overline{\mathbf{S}} \nabla \ln \overline{\rho}) \quad (23)$$

is the total (turbulent and microscopic) viscous force and  $\mathbf{S}$  is given by Eq. (3).

We consider two- and three-dimensional models of an isentropically stratified atmosphere, where the gravitational potential is written as  $\Phi(z) = (z - z_\infty)g$ , so the gravity vector  $\mathbf{g} = -\nabla \Phi = (0, 0, -g)$  is constant. We arrange the initial profiles of density and sound speed such that they take given reference values at  $z = 0$ , i.e.  $\overline{\rho} = \rho_0$  and  $c_s = c_{s0}$  at  $z = 0$ . This implies that  $z_\infty = (3/2)c_{s0}^2/g$ . Our initial profiles therefore obey

$$\rho/\rho_0 = (c_s/c_{s0})^3, \quad c_s^2 = -\frac{2}{3}\Phi. \quad (24)$$



**Fig. 2** Density stratification in a model with  $z_{\text{top}} = H_{\rho 0}$ . The dashed lines indicate the reference value with  $\rho = \rho_0$  at  $z = 0$ .

The local density scale height is  $H_\rho = c_s^2/g$ , and the value of  $H_\rho$  at  $z = 0$  is  $H_{\rho 0} = c_{s0}^2/g$ . In Fig. 2 we show the vertical dependence of the initial density.

We allow for the presence of an imposed field in the  $y$  direction,  $\mathbf{B}_0 = (0, B_0, 0)$ . The total field is then written as  $\overline{\mathbf{B}} = \mathbf{B}_0 + \nabla \times \overline{\mathbf{A}}$ , (25)

so the departure from the imposed field is expressed in terms of the mean magnetic vector potential  $\overline{\mathbf{A}}$ .

On the upper and lower boundaries we adopt stress-free boundary conditions for velocity, i.e.  $\overline{U}_{x,z} = \overline{U}_{y,z} = \overline{U}_z = 0$ , and a normal-field condition for the magnetic field, i.e.  $\overline{B}_x = \overline{B}_y - B_0 = 0$ , corresponding to  $\overline{A}_{x,z} = \overline{A}_{y,z} = \overline{A}_z = 0$  for the vector potential. Here, commas denote partial differentiation. No boundary condition for the density is required. Again, all computations have been carried out with the PENCIL CODE.

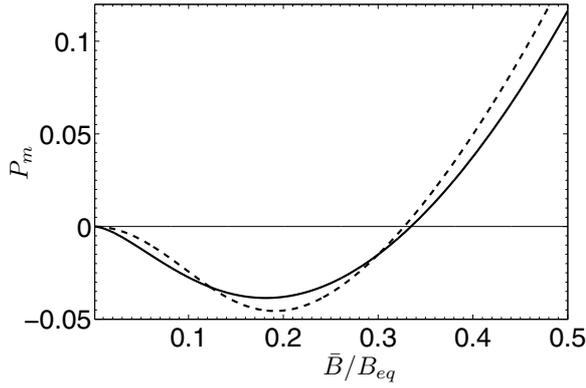
In setting up our model we introduce the wavenumber  $k_f$  of the energy-carrying eddies. This relates the turbulent magnetic diffusivity to the rms velocity via  $\eta_t = u_{\text{rms}}/3k_f$ . This means that the ratio of our non-dimensional field strength to turbulent diffusivity, i.e.  $B_0/c_{s0}\rho_0^{1/2}$  to  $\eta_t/c_{s0}H_\rho$ , is given by  $3k_f H_0$  times  $B_0/B_{\text{eq}}$ .

In this paper we approximate  $q_p$  and  $q_s$  by simple profile functions,

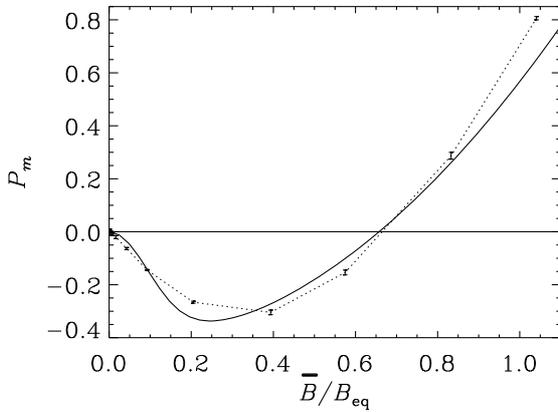
$$q_p = q_{p0} \left( 1 - \frac{2}{\pi} \arctan \frac{\overline{\mathbf{B}}^2}{B_p^2} \right), \quad (26)$$

$$q_s = q_{s0} \left( 1 - \frac{2}{\pi} \arctan \frac{\overline{\mathbf{B}}^2}{B_s^2} \right). \quad (27)$$

Following Rogachevskii & Kleeorin (2007), we also define  $Q_p = 1 - q_p$  and  $Q_s = 1 - q_s$ . Correspondingly, we define the coefficients  $Q_{p0} = 1 - q_{p0}$  and  $Q_{s0} = 1 - q_{s0}$ . For our fiducial model we use  $Q_{p0} = 2Q_{s0} = 20$ , corresponding to  $q_{p0} = 21$  and  $q_{s0} = 11$ . The resulting magnetic pressure is shown in Fig. 3 for their analytic theory and the result from DNS is shown in Fig. 4, together with the corresponding fits. For our fiducial model we choose furthermore



**Fig. 3** The effective mean magnetic pressure  $P_m(\bar{B}) = (1 - q_p)\bar{B}^2/\bar{B}_p^2$  determined by Rogachevskii & Kleeorin (2007) – solid line, and by the model described by Eq. (26) – dashed line ( $\bar{B}_p = 0.21 c_{s0}\rho_0^{1/2}$  and  $q_{p0} = 4$ ).



**Fig. 4** Same as Fig. 3, but from simulation (dotted line). The solid line shows a fit [Eq. (26)] with  $\bar{B}_p = 0.022 c_{s0}\rho_0^{1/2}$  (corresponding to  $\bar{B}_p/B_{eq} = 0.18$ ) and  $q_{p0} = 21$ .

$\bar{B}_p = \bar{B}_s = 0.1 c_{s0}\rho_0^{1/2}$ . For the imposed field strength we choose  $B_0/c_{s0}\rho_0^{1/2} = 0.01$  and for the turbulent magnetic diffusivity we take  $\eta_t/c_{s0}H_\rho = 0.01$ . As discussed above, this means that  $B_0/B_{eq} = 1/3$  if we assume  $k_f H_0 = 1$ .

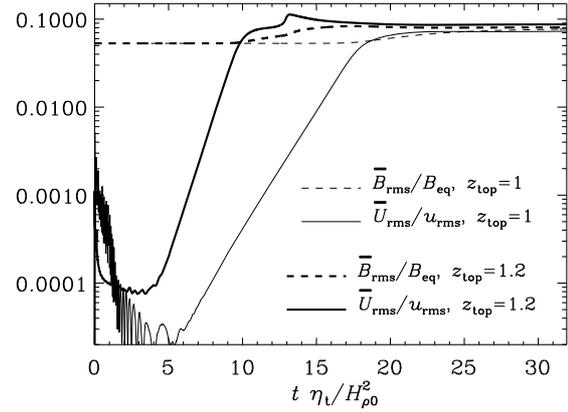
### 3.1 Magnetic structures in two-dimensions

In this paper we consider both two-dimensional and three-dimensional solutions. We begin with two-dimensional models in the  $xz$  plane with an imposed field in the normal ( $y$ ) direction,  $\mathbf{B}_0 = (0, B_0, 0)$ .

In Table 1 we present a summary of some exploratory runs where we list the nondimensional growth rate  $\tilde{\lambda} \equiv \lambda H_\rho^2/\eta_t$ , as well as the nondimensional saturation values of rms mean velocity and mean magnetic field,  $\tilde{U}_{rms} \equiv \bar{U}_{rms}H_\rho/\eta_t$  and  $\tilde{B}_{rms} \equiv \bar{B}_{rms}H_\rho/\eta_t\rho_0^{1/2}$ . These experiments show that decreasing  $Q_{p0}$  and  $Q_{s0}$  lowers the growth rate and the saturation values, while increasing the degree of stratification (e.g., increasing  $\tilde{z}_{top} \equiv z_{top}/H_\rho$  from 1 to

**Table 1** Summary of non-dimensional run parameters together with the resulting non-dimensional growth rates  $\tilde{\lambda}$  as well as the non-dimensional saturation values of rms velocity and magnetic field. The tildes indicate non-dimensional quantities, as explained in Sect. 3.1. Our fiducial run is Run C.

	$\tilde{B}_0$	$\tilde{\eta}_t$	$\tilde{z}_{top}$	$Q_{p0}$	$Q_{s0}$	$\tilde{\lambda}$	$\tilde{U}_{sat}$	$\tilde{B}_{sat}$
A	0.01	0.01	1	-4	-2	0.0005	0.000	0.000
B	0.01	0.01	1.2	-20	-10	0.012	0.013	0.013
C	0.01	0.01	1	-20	-10	0.006	0.013	0.013
D	0.02	0.01	1	-20	-10	0.032	0.025	0.026
E	0.01	0.02	1	-20	-10	0.001	0.001	0.010



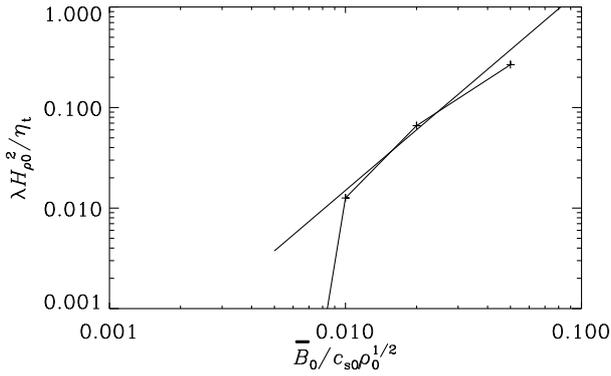
**Fig. 5** Growth of the rms value of mean velocity and mean magnetic field for two runs with different degree of stratification.

1.2, corresponding to an increase of bottom to top density ratio from 108 to 237) enhances them. Likewise, increasing the strength of the imposed field enhances the growth rate and the saturation values, while increasing the magnetic diffusivity lowers them. These results are in agreement with Eq. (19).

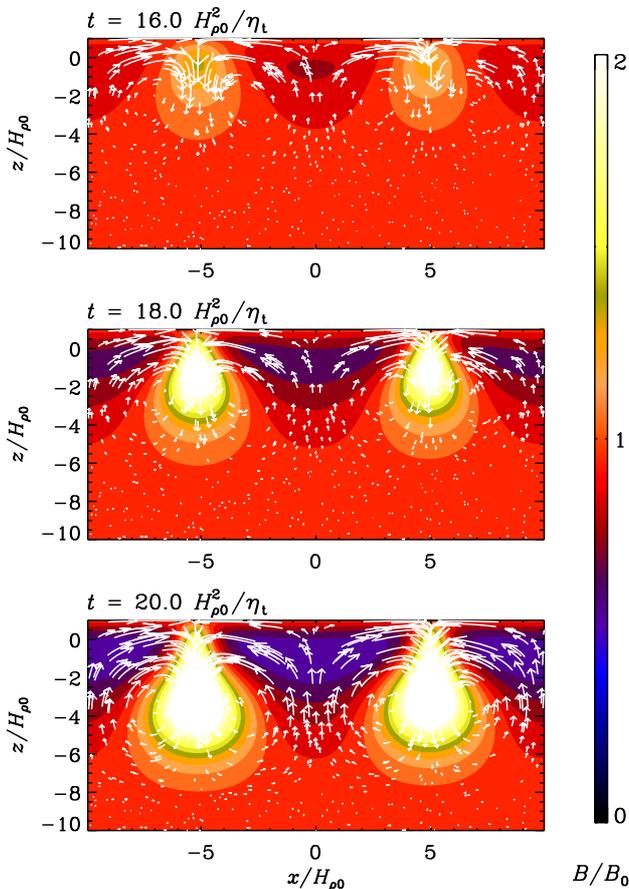
In Fig. 5 we compare the evolution of the rms values of velocity and magnetic field for two different stratifications. As discussed above, increasing the amount of stratification increases the growth rate. The scaling of the growth rate with the strength of the imposed field is shown in Fig. 6.

Next, we consider the structure of the resulting magnetic field. In Fig. 7 we show “meridional” ( $x, z$ ) cross-sections of the magnetic field at three different times during the early phase where the magnetic field just begins to saturate. The horizontal wavelength of the pattern is about  $10 H_\rho$ . As time goes on, the structures of enhanced magnetic field intensify. The decrease of the effective pressure makes them even heavier which explains their subsequent descent.

At later times new structures can form near the surface. In Fig. 8 we show an example during the fully saturated phase of the instability, where one sees the emergence of a new patch that is then swept to the larger one and merges.



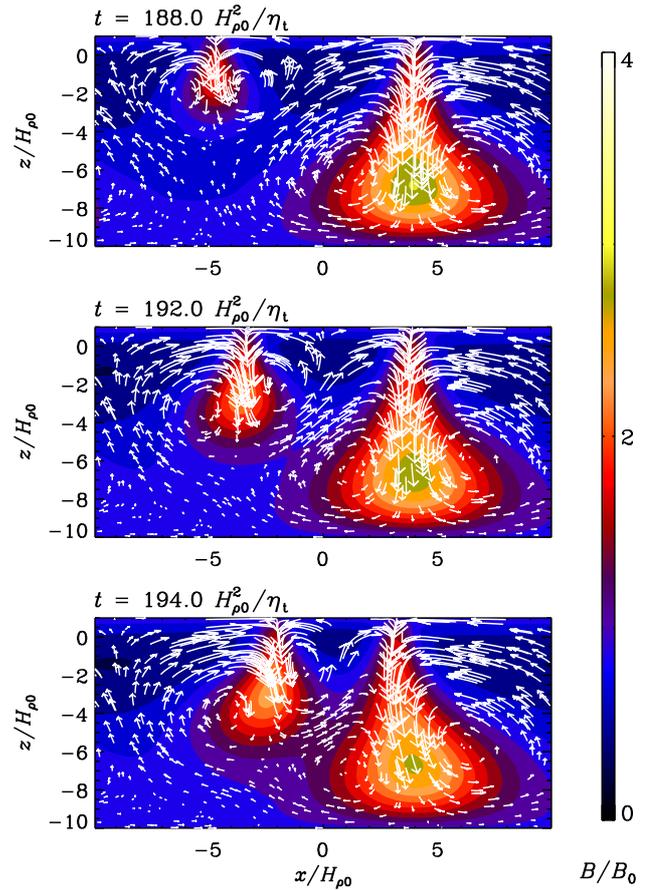
**Fig. 6** Growth rate as a function of  $B_0$ , keeping all other parameters as for the fiducial Run C.



**Fig. 7** (online colour at: [www.an-journal.org](http://www.an-journal.org)) Early evolution of magnetic field in the  $y$  direction (color coded) together with velocity vectors in the  $x, z$  plane. Note the spontaneous production of flux structures.

### 3.2 Magnetic structures in three-dimensions

In order to demonstrate the three-dimensional nature of the instability we extend the domain in the  $y$  direction, so both  $x$  and  $y$  are between  $\pm 10H_{\rho_0}$ , and  $-10 \leq z/H_{\rho_0} \leq 1$ . These simulations are otherwise similar to those in the two-



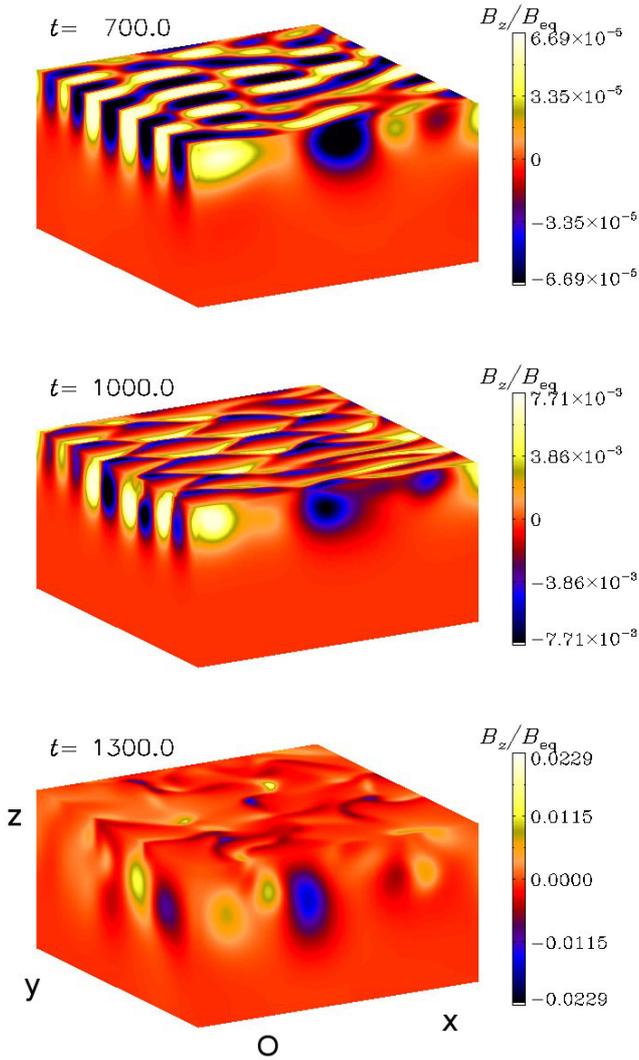
**Fig. 8** (online colour at: [www.an-journal.org](http://www.an-journal.org)) Later evolution of magnetic field for the same run as in Fig. 7. Note the mutual merging of flux structures.

dimensional cases. In Fig. 9 we show visualizations of the field at three characteristic times. Note in particular that the wavelength of the pattern in the  $y$  direction (parallel to the field) is 3–4 times shorter than that in the  $x$  direction (perpendicular to the field). Again, the instability begins to emerge first near the surface where it develops magnetic structures which begin to sink downwards. Viewed from above, one sees the emergence of what looks like multiple bipolar regions.

In Fig. 10 we show a horizontal cross-section from another simulation, where the horizontal extent is only half as much as before. This figure is suggestive of the formation of bipolar structures. The plane has been rotated by  $90^\circ$  such that the  $y$  direction points now from left to right. The black horizontal bar gives the density scale height at the depth of the cross-section, which is about twice the value  $H_{\rho_0}$  at the reference depth.

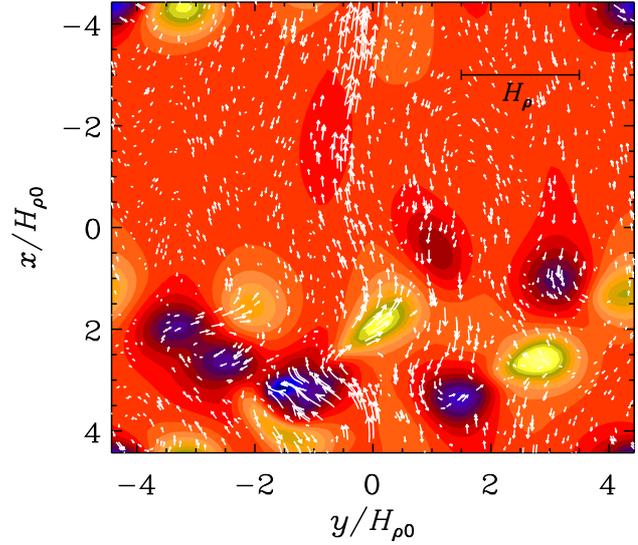
## 4 Conclusions

In this study we have demonstrated in DNS the effects of turbulence on the mean Lorentz force. This effect is quantified by determining the relevant functions  $q_p(\overline{\mathbf{B}})$  and  $q_s(\overline{\mathbf{B}})$



**Fig. 9** (online colour at: [www.an-journal.org](http://www.an-journal.org)) Visualizations of the magnetic field at the early saturation phase ( $t = 700 H_{\rho 0}/c_{s0} = 7 H_{\rho 0}^2/\eta_t$ , *top*), at an intermediate time ( $t = 1000 H_{\rho 0}/c_{s0} = 10 H_{\rho 0}^2/\eta_t$ , *middle*), and a later time ( $t = 1300 H_{\rho 0}/c_{s0} = 13 H_{\rho 0}^2/\eta_t$ , *bottom*). Note the broad similarity of field in the  $x, z$  plane with the two-dimensional cases. Note that the wavelength of the pattern is shorter in the  $y$  direction than in the  $x$  direction. In the final time much of the magnetic field structures have sunk beneath the surface, leaving only a few isolated bipolar structures at the surface.

that relate the sum of the turbulent Reynolds and Maxwell stresses with the Maxwell stress of the mean magnetic field. Using three-dimensional simulations of forced hydromagnetic turbulence with an imposed field, we confirm that the function  $q_p(\overline{\mathbf{B}})$  is positive and can reach values much larger than unity for  $\overline{\mathbf{B}}/B_{\text{eq}} \ll 1$ . This thereby reverses the sign of the effective magnetic pressure  $P_{\text{eff}}(\overline{\mathbf{B}})$  associated with the mean magnetic field, which then becomes  $P_{\text{eff}}(\overline{\mathbf{B}}) = \frac{1}{2}(1 - q_p)\overline{\mathbf{B}}^2$ . We find that the function  $q_s(\overline{\mathbf{B}})$  that determines the modification of magnetic tension, also



**Fig. 10** (online colour at: [www.an-journal.org](http://www.an-journal.org)) Magnetic field in the  $x, y$  plane. Note that the plane has been rotated by  $90^\circ$  in the counterclockwise direction, so the  $y$  direction, corresponding to the azimuthal direction when applied to the Sun, points from left to right. The black horizontal bar gives the density scale height at the depth of the cross-section, which is about twice the value  $H_{\rho 0}$  at the reference depth.

reaches values much larger than unity, but its value is less than half the value of  $q_p$  and the error bars are larger.

This work has also demonstrated explicitly the possibility of a large-scale instability of the full system of mean-field MHD equations. Our study has revealed the presence of spatial structures arising from the instability that might be associated with the formation of bipolar magnetic regions and perhaps also sunspots when applied to the Sun.

The effects of turbulence on the large-scale Lorentz force may also be important in applications to the solar torsional oscillations by changing the effective coefficient in front of the magnetic tension force,  $\overline{\mathbf{B}} \cdot \nabla \overline{\mathbf{B}}$  (Rüdiger et al. 1986; Rüdiger & Kichatinov 1990). More specifically, these turbulence effects may be critical for explaining the narrow structure of the observed solar torsional oscillations (Kleeorin et al. 1996).

Clearly, there are several possibilities of improvement that might make the model more realistic and eventually suitable for confrontation against observations. Most important is perhaps the fact that our reference values  $B_{p0}$  and  $B_{s0}$  in the quenching profiles (26) and (27), as well as the coefficients  $q_{p0}$  and  $q_{s0}$  in these profile functions, are currently kept constant. It will be more realistic to make them vary with depth, because density and turbulent rms velocity vary with depth. Another extension of the model would be to allow for the possibility that the magnetic field is generated by a mean-field dynamo rather than relying on an imposed field.

On the more theoretical side, there is a need for further verification of the essential physics of the negative magnetic pressure effect. In particular, one must wonder why the effects of this instability have not yet been seen in DNS. A likely possibility is the lack of sufficient scale separation. Only now realistic simulations are beginning to be large enough to encompass sufficiently many cells in all three directions. The importance of sufficient horizontal extent was already emphasized by Tao et al. (1998), who find clear evidence of a segregation into strongly magnetized and weakly magnetized regions, a phenomenon that might be closely related to that reported here. A particularly promising approach might be to generalize the direct simulations discussed in the present paper to the case with vertical density stratification such that the setup becomes similar to the mean-field models that we also studied in this paper. The turbulence would then still be driven by a forcing function. Another possibility is to study this effect with turbulent convection instead of forced turbulence. This is particularly interesting, because theoretical predictions of Rogachevskii & Kleeorin (2007) suggest that the modification of the effective Lorentz force will be even stronger in turbulent convection.

*Acknowledgements.* We thank the referee for critical remarks. We have benefited from very useful discussions with Alexander Kosovichev. We acknowledge the use of computing time at the Center for Parallel Computers at the Royal Institute of Technology in Sweden as well as the National Supercomputing Center in Linköping. This work was supported in part by the European Research Council under the AstroDyn project 227952 and the Swedish Research Council grant 621-2007-4064. NK and IR thank NORDITA for hospitality and support during their visits. The final stage of this work was completed while participating at the NORDITA program on “Solar and stellar dynamos and cycles”.

## References

- Brandenburg, A.: 2005, *ApJ* 625, 539  
 Choudhuri, A.R., D’Silva, S.: 1990, *A&A* 239, 326  
 D’Silva, S., Choudhuri, A.R.: 1993, *A&A* 272, 621  
 Ferriz-Mas, A., Schüssler, M.: 1993, *GApFD* 72, 209  
 Gilman, P.A.: 1970a, *ApJ* 162, 1019  
 Gilman P.A.: 1970b, *A&A* 286, 305  
 Haugen, N.E.L., Brandenburg, A.: 2004, *Phys Rev E* 70, 036408  
 Haugen, N.E.L., Brandenburg, A., Dobler, W.: 2004, *Phys Rev E* 70, 016308  
 Heinemann, T., Nordlund, Å., Scharmer, G.B., Spruit, H.C.: 2007, *ApJ* 669, 1390  
 Hindman, B.W., Haber, D.A., Toomre, J.: 2009, *ApJ* 698, 1749  
 Hurlburt, N.E., Rucklidge, A.M.: 2000, *MNRAS* 314, 793  
 Kitchatinov, L.L., Mazur, M.V.: 2000, *SoPh* 191, 325  
 Kitiashvili, I.N., Kosovichev, A.G., Wray, A.A., Mansour, N.N.: 2009, *ApJ* 700, 178  
 Kleeorin, N., Mond, M., Rogachevskii, I.: 1996, *A&A* 307, 293  
 Kleeorin, N., Rogachevskii, I.: 1994, *Phys Rev E* 50, 2716  
 Kleeorin, N.I., Rogachevskii, I.V., Ruzmaikin, A.A.: 1989, *SvAL* 15, 274  
 Kleeorin, N.I., Rogachevskii, I.V., Ruzmaikin, A.A.: 1990, *SJETP* 70, 878  
 Kosovichev, A.G.: 2002, *AN* 323, 186  
 Krause, F., Rädler, K.-H.: 1980, *Mean-Field Magnetohydrodynamics and Dynamo Theory*, Pergamon Press, Oxford  
 Landau, L.D., Lifshitz, E.M.: 1975, *Classical Theory of Fields*, Pergamon, Oxford  
 Landau, L.D., Lifshitz, E.M.: 1984, *Theory of Elasticity*, Pergamon, Oxford  
 Moffatt, H.K.: 1978, *Magnetic Field Generation in Electrically Conducting Fluids*, Cambridge University Press, Cambridge  
 Moreno-Insertis, F.: 1986, *A&A* 166, 291  
 Newcomb, W.A.: 1961, *PhFl* 4, 391  
 Parker, E.N.: 1955, *ApJ* 121, 491  
 Parker, E.N.: 1966, *ApJ* 145, 811  
 Parker, E.N.: 1979a, *ApJ* 230, 905  
 Parker, E.N.: 1979b, *Cosmical Magnetic Fields*, Oxford University Press, New York  
 Parker, E.N.: 1982, *ApJ* 256, 302  
 Parker, E.N.: 1984, *ApJ* 283, 343  
 Priest, E.R.: 1982, *Solar Magnetohydrodynamics*, D. Reidel Publ. Co., Dordrecht  
 Rädler, K.-H.: 1974, *AN* 295, 265  
 Rempel, M., Schüssler, M., Cameron, R.H., Knölker, M.: 2009, *Sci* 325, 171  
 Rogachevskii, I., Kleeorin, N.: 2007, *Phys Rev E* 76, 056307  
 Rüdiger, G.: 1974, *AN* 295, 275  
 Rüdiger, G.: 1980, *GApFD* 16, 239  
 Rüdiger, G.: 1989, *Differential Rotation and Stellar Convection: Sun and Solar-Type Stars*, Gordon & Breach, New York  
 Rüdiger, G., Hollerbach, R.: 2004, *The Magnetic Universe* Wiley-VCH, Weinheim  
 Rüdiger, G., Kichatinov, L.L.: 1990, *A&A* 236, 503  
 Rüdiger, G., Tuominen, I., Krause, F., Virtanen, H.: 1986, *A&A* 166, 306  
 Schatten, K.H.: 2009, *SoPh* 255, 3  
 Schekochihin, A.A., Haugen, N.E.L., Brandenburg, A., Cowley, S.C., Maron, J.L., McWilliams, J.C.: 2005, *ApJ* 625, L115  
 Schekochihin, A.A., Iskakov, A.B., Cowley, S.C., McWilliams, J.C., Proctor, M.R.E., Yousef, T.A.: 2007, *NJPh* 9, 300  
 Schüssler, M., Caligari P., Ferriz-Mas A., Moreno-Insertis, F.: 1994, *A&A* 281, L69  
 Spiegel, E.A., Weiss, N.O.: 1980, *Nature* 287, 616  
 Spruit, H.C. 1981, *A&A* 98, 155  
 Spruit, H.C., van Ballegoijen, A.A.: 1982, *A&A* 106, 58  
 Tao, L., Weiss, N.O., Brownjohn, D.P., Proctor, M.R.E.: 1998, *ApJ* 496, L39  
 Thomas, J.H., Weiss, N.O., Tobias, S.M., Brummell, N.H.: 2002, *Nature* 420, 390  
 Tobias, S.M., Hughes, D.W.: 2004, *ApJ* 603, 785  
 Tobias, S.M., Brummell, N.H., Clune T.L., Toomre, J.: 2001, *ApJ* 549, 1183  
 Zhao, J., Kosovichev, A.G.: 2004, *ApJ* 603, 776  
 Zhao, J., Kosovichev, A.G., Duvall, T.: 2001, *ApJ* 557, 384